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## Magellan/M2FS Spectroscopy of the Reticulum 2 Dwarf Spheroidal Galaxy

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# MAGELLAN/M2FS SPECTROSCOPY OF THE RETICULUM 2 DWARF SPHEROIDAL GALAXY

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## ABSTRACT

We present results from spectroscopic observations with the Michigan/Magellan Fiber System (M2FS) of 182 stellar targets along the line of sight to the newly-discovered ‘ultrafaint’ object Reticulum 2 (Ret 2). For 38 of these targets, the spectra are sufficient to provide simultaneous estimates of line-of-sight velocity ( $v_{\text{los}}$ , median random error  $\delta v_{\text{los}} = 1.3 \text{ km s}^{-1}$ ), effective temperature ( $T_{\text{eff}}$ ,  $\delta T_{\text{eff}} = 464 \text{ K}$ ), surface gravity ( $\log g$ ,  $\delta \log g = 0.54 \text{ dex}$ ) and iron abundance ( $[\text{Fe}/\text{H}]$ ,  $\delta [\text{Fe}/\text{H}] = 0.45 \text{ dex}$ ). We use these results to confirm 18 stars as members of Ret 2. From the member sample we estimate a velocity dispersion of  $\sigma_{v_{\text{los}}} = 3.6^{+0.9}_{-0.6} \text{ km s}^{-1}$  about a mean of  $\langle v_{\text{los}} \rangle = 64.8^{+1.1}_{-1.0} \text{ km s}^{-1}$  in the solar rest frame ( $\sim -90.9 \text{ km s}^{-1}$  in the Galactic rest frame), and a metallicity dispersion of  $\sigma_{[\text{Fe}/\text{H}]} = 0.50^{+0.17}_{-0.13} \text{ dex}$  about a mean of  $\langle [\text{Fe}/\text{H}] \rangle = -2.67^{+0.34}_{-0.34}$ . These estimates marginalize over possible velocity and metallicity gradients, which are consistent with zero. Our results place Ret 2 on chemodynamical scaling relations followed by the Milky Way’s dwarf-galactic satellites. Under assumptions of dynamic equilibrium and negligible contamination from binary stars—both of which must be checked with deeper imaging and repeat spectroscopic observations—the estimated velocity dispersion suggests a dynamical mass of  $M(R_h) \approx 5R_h\sigma_{v_{\text{los}}}^2/(2G) = 2.4^{+1.3}_{-0.8} \times 10^5 M_{\odot}$  enclosed within projected half-light radius  $R_h \sim 32 \text{ pc}$ , with mass-to-light ratio  $\approx 2M(R_h)/L_V = 462^{+264}_{-157}$  in solar units.

*Subject headings:* galaxies: dwarf — galaxies: individual (Reticulum 2) — (galaxies:) Local Group — galaxies: kinematics and dynamics — methods: data analysis — techniques: spectroscopic

## 1. INTRODUCTION

The census of Local Group galaxies has been revised dramatically and repeatedly over the last decade. Mining of the SDSS stellar catalog has yielded discoveries of  $\sim 15$  low-luminosity, dwarf-galactic satellites of the Milky Way (e.g., Willman et al. 2005; Zucker et al. 2006; Belokurov et al. 2007). The PanDAS and PanStarrs surveys have found nearly two dozen new satellites of M31 (McConnachie et al. 2009; Martin et al. 2013). All told, the population of known Local Group galaxies has nearly tripled (McConnachie 2012). Of the new members, perhaps the most intriguing are the Milky Way’s ‘ultrafaint’ satellites. These objects have lowered the floor of the observed galaxy luminosity function from  $M_V \sim -8$  to  $M_V \sim -2$ , such that some galaxies’ total luminosities can be dominated by a single red giant star (Martin et al. 2008). Moreover, the structural parameters of ultrafaints have blurred what was once an obvious distinction between the Milky Way’s dwarf-galactic satellites and its globular clusters. As a result, the proper classification of most ultrafaint objects now requires spectroscopic measurements of velocity dispersions, metallicities and metallicity dispersions that can indicate the presence of a dark matter halo.

Most recently, the Dark Energy Survey (DES) has revealed nine new Galactic satellites at southern latitudes (Koposov et al. 2015; The DES Collaboration et al. 2015, ‘K15’ and

‘DES15’ hereafter). Seven of the new objects have sizes and luminosities characteristic of ultrafaints. One of them, Reticulum 2 (Ret 2), has already attracted attention for several reasons. First, Ret 2 is the nearest ( $D \sim 30 \text{ kpc}$ ) and most easily detected of the newly-discovered objects. Second, Ret 2 clearly has a flattened morphology, which may indicate ongoing tidal disruption or perhaps rotation. Third, using public data from the Fermi-LAT, Geringer-Sameth et al. (2015) find evidence for gamma-ray emission that is consistent with dark matter annihilation in Ret 2. The Fermi-LAT Collaboration et al. (2015) assign low significance to the gamma-ray signal based on unreleased Fermi-LAT data; however, in an independent analysis of the public data, Hooper & Linden (2015) reproduce the original detection. In any case, Ret 2 is clearly an intriguing object, and the first question to settle is whether Ret 2 presents chemo-dynamical evidence for dark matter. That is, is Ret 2 a globular cluster or a galaxy?

Here we present results from an initial spectroscopic ‘re-connaissance’ of individual stellar targets along the line of sight to Ret 2. We identify a sample consisting of 18 member stars, which we use to characterize Ret 2’s chemodynamical properties. Specifically, we estimate the means and dispersions of velocity and metallicity distributions, and we check for velocity and metallicity gradients that might provide clues about Ret 2’s dynamical state and formation history. Finally, we compare Ret 2’s properties to those of known dwarf galaxies and globular clusters in order to determine which population can claim Ret 2 as its newest member.

## 2. OBSERVATIONS AND DATA REDUCTION

On 19 February 2015, we observed Ret 2 with the Michigan/Magellan Fiber System (M2FS, Mateo et al. 2012) on the 6.5-m Magellan/Clay Telescope at Las Campanas Observatory, Chile. M2FS deploys up to 256 optical fibers—each with

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an entrance aperture of diameter 1.2 arcsec—anywhere over a (f/11) focal surface with diameter 29 arcmin. The fibers feed twin spectrographs that can be configured independently in a wide variety of modes. For the Ret 2 observations, both M2FS spectrographs were configured identically in ‘HiRes’ mode, with order-isolation filters providing coverage over the range 5132–5186 Å at effective resolution  $\mathcal{R} \sim 18000$ .

The left-hand panel of Figure 1 shows a color-magnitude diagram for stars projected within 15 arcmin of Ret 2’s center ( $\alpha_{2000} = 53.9256$  deg,  $\delta_{2000} = -54.0492$  deg), from the photometric catalog derived by K15 in their analysis of public images from the Dark Energy Survey. With M2FS we targeted 182 stars (large circles in Figure 1) selected as probable members of Ret 2 by the DES’s Milky Way Science Working Group, who also provided the coordinates necessary for designing M2FS plug plates (private communication, 2015). The right-hand panel of Figure 1 shows positions of the targeted stars. We also allocated 64 fibers to regions of blank sky in order to estimate background. We observed the Ret 2 field for a total of 2 hours ( $3 \times 2400$  s). At evening twilight the field was already rather low in the west, such that airmass during our observations increased from 1.3 to 1.7 and seeing hovered above 1.2 arcsec.

We analyze the individual images (i.e. the single 40-min exposures) as well as the ‘stacked’ image consisting of the average over the three individual images. Other raw images include ancillary exposures of emission-line sources (‘arc’ spectra), continuum sources (‘quartz’ spectra), twilight exposures, bias and dark exposures, all obtained using the same configuration and detector binning ( $2 \times 2$ ) as the science observations.

We process the raw data using standard IRAF routines, following procedures described in detail by Mateo et al. (2015). To summarize, we begin by performing overscan, bias and dark current corrections. We then average the three individual exposures to produce a single, stacked science frame for each of the two spectrographs. From science frames corresponding to individual as well as stacked exposures, we extract one-dimensional spectra. For a given spectrum we obtain the wavelength solution by fitting a 4<sup>th</sup>-order cubic spline to  $\sim 30$  emission lines identified in the identically-extracted arc spectrum from the same aperture. Residuals to the wavelength solution typically have rms  $\sim 0.3$  km s<sup>-1</sup>. We correct for variations in fiber throughput by dividing each spectrum by a normalized fit to the continuum in the twilight spectrum obtained in the same aperture. Finally, we estimate and subtract the mean contribution from sky background following the procedure introduced by Koposov et al. (2011), which avoids re-binning science spectra.

In parallel, we also compute variance spectra that account for all known sources of noise in the processed science spectra. These include Poisson noise, readout noise, rms fluctuations in the bias and dark images, and background noise propagated from individual sky spectra into the mean sky spectrum. We assign large variances ( $10^{100}$ ) to pixels affected by cosmic rays, which we identify as outliers above an iterative fit to the continuum in the science spectrum.

Figure 2 displays examples of sky-subtracted M2FS spectra (from stacked science frames) for nine representative Ret 2 targets. Those in left-hand panels are from likely members of Ret 2 (Section 4.1) and show the relatively weak absorption features that are typical of metal-poor subgiants. Spectra in

upper right-hand panels belong to likely foreground contaminants and show the stronger absorption characteristic of late-type dwarfs. Lower right-hand panels display spectra for two stars identified as blue horizontal branch candidates by both K15 and DES15. These are the two bluest stars in our sample, and while their continua are well sampled, they exhibit no obvious absorption features in the observed spectral range (see Section 3.5). Finally, Figure 3 shows two anomalous spectra that reveal their sources to be things other than single stars. We discuss these spectra in more detail in Section 3.5.

### 3. MODELLING OF M2FS SPECTRA

We model the sky-subtracted M2FS spectra following the procedure of Walker et al. (2015, ‘W15’ hereafter). Briefly, we adopt a spectral model of the form (see also Koleva et al. 2009; Koposov et al. 2011)

$$M(\lambda) = P_l(\lambda) T \left( \lambda \left[ 1 + \frac{Q_m(\lambda) + v_{\text{los}}}{c} \right] \right), \quad (1)$$

where  $c$  is the speed of light, the polynomial  $P_l(\lambda)$  gives shape to the continuum and  $T \left( \lambda \left[ 1 + \frac{Q_m(\lambda) + v_{\text{los}}}{c} \right] \right)$  is a continuum-normalized template spectrum. The template is redshifted according to line-of-sight velocity,  $v_{\text{los}}$ , and also by a second polynomial,  $Q_m(\lambda)$ , that adjusts for systematic differences between wavelength solutions of real and template spectra. We generate templates from a synthetic library that has previously been used for the SEGUE Stellar Parameter Pipeline (‘SSPP’, Lee et al. 2008a,b) and which is calculated over a regular grid in effective temperature ( $T_{\text{eff}}$ ), surface gravity ( $\log g$ ) and metallicity ([Fe/H]).

Including nuisance parameters that specify polynomial coefficients and smooth the templates according to the instrumental line-spread function (LSF), the full model has 15 free parameters. Four have physical meaning:  $v_{\text{los}}$ ,  $T_{\text{eff}}$ ,  $\log g$  and [Fe/H]. In order to obtain simultaneous estimates of all parameters, we follow W15’s Bayesian analysis. We adopt the same likelihood function given by W15’s equation 9, and the same priors listed in W15’s Table 2, with one exception: for  $h_0$ , the parameter that specifies the amount by which the template is smoothed in order to mimic the instrumental LSF, we adopt a prior that is uniform between 0.001 Å and 0.1 Å (W15’s prior on  $h_0$  is uniform between 0.05 Å and 0.1 Å). We find that this adjustment improves fits to the narrowest absorption features.

#### 3.1. Posterior PDFs

We use the software package MultiNest (Feroz & Hobson 2008; Feroz et al. 2009) to scan the parameter space and generate random samplings of the 15-dimensional posterior PDF. For a given parameter,  $X$ , we then calculate moments of the marginalized, 1D posterior PDF as follows. The first moment is the mean,  $\bar{X} \equiv N^{-1} \sum_{i=1}^N x_i$ , which we take to be the central value. The second moment is the variance,  $\sigma_X^2 \equiv (N-1)^{-1} \sum_{i=1}^N (x_i - \bar{X})^2$ , which we take to be the square of the  $1\sigma$  credibility interval. The third moment is skewness,  $S \equiv N^{-1} \sum_{i=1}^N [(x_i - \bar{X}) / \sqrt{\sigma_X^2}]^3$ , which equals zero for a symmetric distribution. The fourth moment is kurtosis,  $K \equiv N^{-1} \sum_{i=1}^N [(x_i - \bar{X}) / \sqrt{\sigma_X^2}]^4$ , which distinguishes Gaussian distributions ( $K = 3$ ) from ‘leptokurtic’ ones with sharper peaks and fatter tails ( $K > 3$ ) and ‘platykurtic’ ones with broader peaks and weaker tails ( $K < 3$ ).

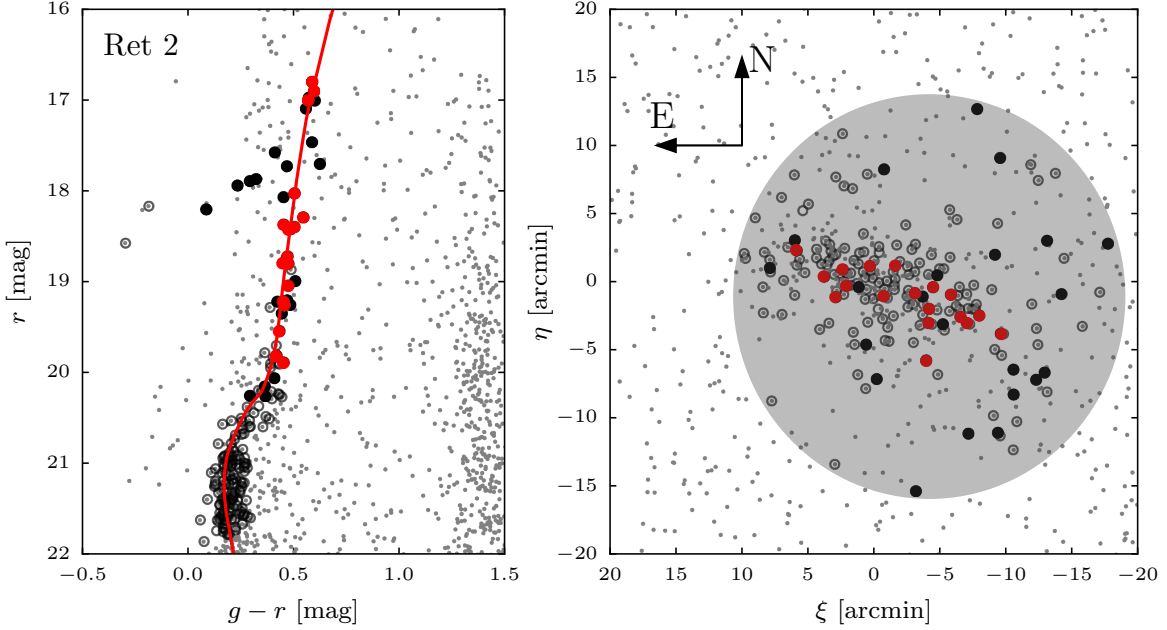


FIG. 1.— *Left*: Color-magnitude diagram for stars within  $R \leq 15$  arcmin of Ret 2’s center (Koposov et al. 2015). The red line is the Dartmouth isochrone for age = 12 Gyr,  $[\text{Fe}/\text{H}] = -2.5$ ,  $[\alpha/\text{Fe}] = +0.4$ , and  $m - M = 17.4$  (Dotter et al. 2008). Large circles enclose stars identified as probable Ret 2 members by the DES MW working group (private communication, 2015) and subsequently targeted by M2FS. Filled circles mark 38 stars whose spectroscopic measurements meet quality-control criteria. Of those, red circles identify stars whose spectroscopic properties are consistent with Ret 2 membership; black circles are spectroscopic non-members. *Right*: Standard coordinates for stars within 0.2 magnitudes of the isochrone shown in the left-hand panel. Markers and their colors have the same meanings as in the left-hand panel. The large shaded circle represents the M2FS field of view.

For all 182 stacked M2FS spectra from Ret 2 targets, Figure 4 shows how moments of PDFs for physical parameters vary with median S/N per pixel. As S/N grows, the PDFs become narrower and more Gaussian ( $S \sim 0$ ,  $K \sim 3$ ). In subsequent analysis we consider only the sample consisting of 38 observations for which the PDF for velocity has  $\sigma \leq 5 \text{ km s}^{-1}$ ,  $|S| \leq 1$  and  $|K - 3| \leq 1$  (red points in Figure 4).

### 3.2. Accuracy and Precision

In order to examine accuracy and precision, we apply the same analysis to 256 solar spectra that we acquired during evening twilight on the same night and with the same spectrograph configuration used for the Ret 2 observations. Reassuringly, parameter estimates from the collection of twilight spectra exhibit empirical scatter that is consistent with variances of the posterior PDFs. Among the 256 twilight spectra, standard deviations of central values for  $v_{\text{los}}$ ,  $T_{\text{eff}}$ ,  $\log g$  and  $[\text{Fe}/\text{H}]$  are  $0.1 \text{ km s}^{-1}$ ,  $43 \text{ K}$ ,  $0.04 \text{ dex}$  and  $0.02 \text{ dex}$ , respectively. Furthermore, mean estimates show only mild systematic offsets with respect to known solar values:  $\langle v_{\text{los}} \rangle - v_{\text{los},\odot} = -3.9 \text{ km s}^{-1}$ ,  $\langle T_{\text{eff}} \rangle - T_{\text{eff},\odot} = 69 \text{ K}$ ,  $\langle \log g \rangle - \log g_{\odot} = 0.09 \text{ dex}$  and  $\langle [\text{Fe}/\text{H}] \rangle - [\text{Fe}/\text{H}]_{\odot} = -0.20 \text{ dex}$ . Following W15, for stellar-atmospheric parameters we treat these offsets as zero-point errors and subtract them from raw estimates obtained for science targets. We also add the empirical scatter, in quadrature, to (square roots of) second moments for all observations. After applying these adjustments, our estimates of stellar-atmospheric parameters have median (minimum, maximum) random errors of  $\sigma_{T_{\text{eff}}} = 464 (50, 1035) \text{ K}$ ,  $\sigma_{\log g} = 0.54 (0.08, 1.30) \text{ dex}$  and  $\sigma_{[\text{Fe}/\text{H}]} = 0.45 (0.04, 1.05) \text{ dex}$ .

For velocities, we do not adjust the zero point based on results from the twilight spectra. In our experience, M2FS velocities shift systematically as ambient temperature changes

rapidly during twilight, and the observed offset between twilight and solar velocities reflects this phenomenon. As pointed out to us by the DES collaboration (private communication, 2015), a temperature-dependent velocity shift continued, albeit at a slower rate, during our observations of Ret 2. In order to examine this effect, we compare velocity estimates that we obtain from fits to spectra obtained in individual science exposures. We consider only the 73 single-exposure spectra whose results satisfy the same quality-control criteria imposed on the stacked spectra (Section 3.1). Consistently with the effect noticed by the DES collaboration, we find that velocity estimates shift systematically over the three exposures and at different rates for the two spectrograph channels. From 39 (34) spectra observed on M2FS’s ‘blue’ (‘red’) channel, we measure a weighted mean difference of  $\langle v_{\text{exp3}} - v_{\text{exp1}} \rangle = 2.1 \pm 0.3 \text{ km s}^{-1}$  ( $1.1 \pm 0.3 \text{ km s}^{-1}$ ) between the first and third exposures. Thus the raw velocities that we estimate from the stacked frames have channel-dependent zero-point errors. In order to compensate for this effect, we subtract the appropriate channel-specific value of  $\langle v_{\text{exp3}} - v_{\text{exp1}} \rangle$  from each of the raw velocities that we estimate from the stacked frames. We also add the corresponding error in our estimate of  $\langle v_{\text{exp3}} - v_{\text{exp1}} \rangle$ , in quadrature, to the error of our velocity estimates, after which our velocity sample has median (minimum, maximum) error  $\sigma_{v_{\text{los}}} = 1.3 (0.4, 3.3) \text{ km s}^{-1}$ . This procedure is designed only to remove the dependence of the velocity zero point on channel, and does not necessarily correct for a global error in zero point. However, a comparison to an independent spectroscopic sample obtained with VLT/FLAMES (Koposov et al., in preparation), shows that the zero points are in agreement. For the 17 stars common to both samples, the mean velocity difference is  $\langle v_{\text{M2FS}} - v_{\text{VLT}} \rangle = 0.3 \pm 0.3 \text{ km s}^{-1}$ .

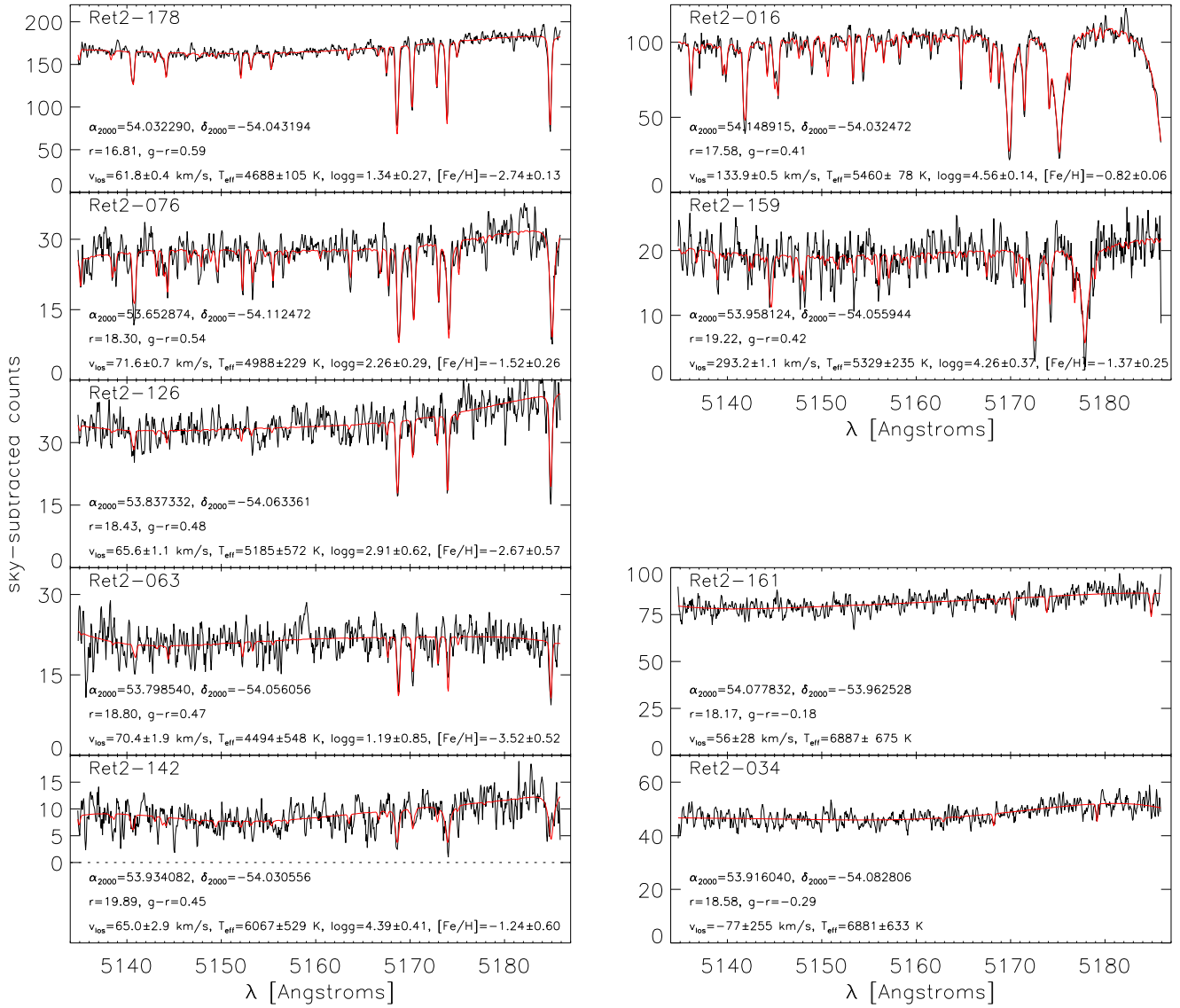


FIG. 2.— Sky-subtracted M2FS spectra, with best-fitting models overplotted. Text gives target ID (upper left), equatorial coordinates (in degrees), color and magnitude, and estimates of redshift and stellar-atmospheric parameters. Panels on the left-hand side show spectra for probable members of Reticulum 2. The two upper-right panels show spectra for probable contaminants in the Galactic foreground. Lower-right panels show spectra for two blue horizontal branch candidates.

### 3.3. Repeatability

In order to gauge repeatability of our estimates for physical parameters, we again consider results from our modelling of 73 single-exposure spectra whose results satisfy quality-control criteria. The top row of panels in Figure 5 compares parameter estimates from the first exposure with those in either of the two later exposures (all velocities from a given exposure and channel are shifted to remove the systematic drift in zero-point discussed above). The bottom row of panels in Figure 5 displays distributions of deviations with respect to weighted means— $\langle \bar{X} \rangle \equiv \sum_{i=1}^N (\bar{X}_i / \sigma_{X_i}^2) / \sum_{i=1}^{N_{\text{obs}}} (1 / \sigma_{X_i}^2)$ —normalized by propagated error. Gray curves show Gaussian distributions with the same integrated area as the histograms. The observed distributions are all narrower than the Gaussian

curves, implying that the variances of the PDFs give credibility intervals that may be slightly overestimated.

As a final check on reliability of our velocity estimates, we independently measure line-of-sight velocities using IRAF’s ‘fxcor’ package. We cross-correlate each target spectrum from the stacked frame with a high-S/N M2FS spectrum that we acquired for the radial velocity standard star CPD-432527 ( $v_{\text{los}} = 19.7 \pm 0.9 \text{ km s}^{-1}$ ; Udry et al. 1999) immediately before the Ret 2 observations. Reassuringly, we find good agreement with results from the Bayesian procedure described above. For the 38 spectra that satisfy quality-control criteria (Section 3.1), the mean deviation between (raw) velocity estimates is  $\sim 0.1 \pm 0.30 \text{ km s}^{-1}$ . Furthermore, when we model the spectrum of CPD-432527 using our Bayesian procedure, we



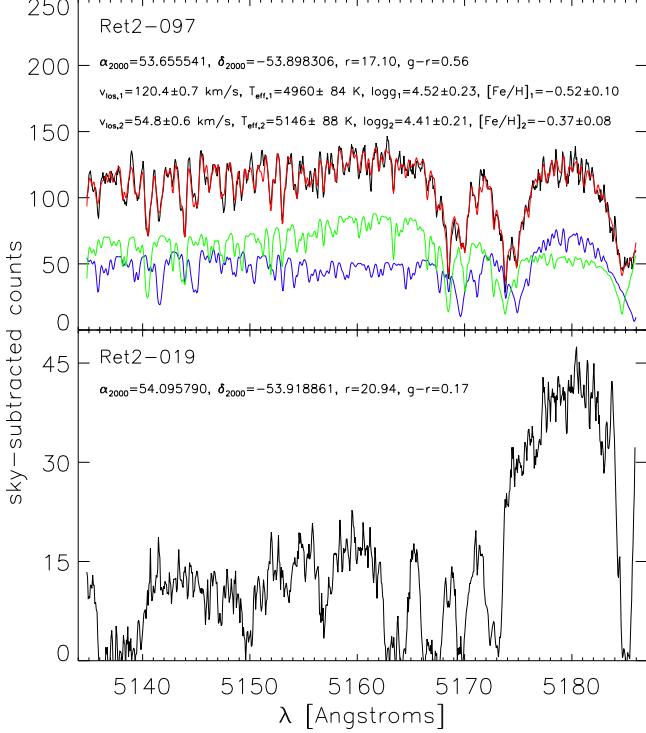


FIG. 3.— Sky-subtracted M2FS spectra for two anomalous sources. *Top:* This source is well-fit by the superposition (red) of two metal-rich model spectra (blue and green). *Bottom:* We speculate that the extremely wide, deep and irregular absorption features correspond to an absorption line system at redshift  $z \gtrsim 2.5$ .

obtain  $v_{\text{los}} = 19.7 \pm 0.3$  km s $^{-1}$ , recovering the previously-published value.

### 3.4. Data

Table 1 lists results for the 38 spectra that satisfy the quality-control criteria discussed in Section 3.1. The online version of this article includes a supplementary table with results for all 182 observed stars. The first three columns of Table 1 list target ID and equatorial coordinates, followed by angular separation from Ret 2’s center and  $g$ ,  $r$  magnitudes from the photometry of K15, corrected for extinction using the dust maps of Schlegel et al. (1998). The seventh column lists median S/N per pixel. Columns 8-11 then list the results of our spectroscopic modelling in terms of the first four moments of PDFs for  $v_{\text{los}}$ ,  $T_{\text{eff}}$ ,  $\log g$  and  $[\text{Fe}/\text{H}]$ . For each parameter, the quoted central value is the first moment (mean) and the credible interval is the square root of the second moment; both values have been adjusted according to the offsets and empirical scatter described in Section 3.2. Skewness and kurtosis are listed parenthetically as superscripts above the credible interval. The last column indicates whether or not the star satisfies membership criteria discussed in Section 4.1. For all observations, the first exposure began at heliocentric Julian date HJD=2457072.542 days and the midpoint of the  $3 \times 40$ -minute exposure occurred at HJD=2457072.579 days.

### 3.5. Special Cases

Four spectra from our sample merit special attention. Table 2 lists their coordinates and  $g$ ,  $r$  magnitudes. The first two (Ret2-161 and Ret2-034) belong to the two blue horizontal branch (BHB) candidates and are displayed in the lower-right

panels of Figure 2. Their spectra show well-sampled continua but no obvious absorption features, as expected given the greater degree of ionization at high temperature. As a result, model parameters for these stars are poorly constrained. While their velocities are loosely consistent with Ret 2 membership (see Section 4.1), the associated uncertainties are  $\sim 30$  km s $^{-1}$  and  $\sim 260$  km s $^{-1}$ , such that neither spectrum satisfies our quality-control criteria (Section 3.1). On the other hand, their temperatures ( $T_{\text{eff}}=6890 \pm 680$  K and  $6880 \pm 630$  K), while also highly uncertain, are the two largest that we obtained for any of our 182 observed spectra. We conclude that these two stars remain strong BHB candidates, with confirmation requiring spectra that cover a larger range in wavelength.

The next special case (Ret2-097) appears in the upper panel of Figure 3. This ‘star’ has color and magnitude placing it near the isochrone and above the horizontal branch in Figure 1. However, the spectrum is rich in broad and apparently double-valleyed absorption features. We model this spectrum as the superposition of *two* spectra of the form given by Equation 1. Allowing for a second set of stellar-atmospheric parameters and a second continuum polynomial, the resulting model has 25 free parameters. As shown in Figure 2, this double-star model gives a reasonably good fit to the spectrum. One of the sources has  $v_{\text{los}} = 54.8 \pm 0.6$  km s $^{-1}$ , slightly below the range of velocities that we attribute to Ret 2 members (Section 4.1). Its strong surface gravity ( $\log g = 4.41 \pm 0.21$ ) and rich metallicity ( $[\text{Fe}/\text{H}] = -0.37 \pm 0.08$ ) are typical of foreground contamination. The other contributing star has  $v_{\text{los}} = 120.4 \pm 0.7$  km s $^{-1}$ ,  $\log g = 4.52 \pm 0.23$  and is also metal-rich, with  $[\text{Fe}/\text{H}] = -0.52 \pm 0.10$ . We speculate that this source is a physical binary composed of two K-type main sequence stars at a distance of a few kpc. Assuming a mass ratio near unity, the center of mass has  $v_{\text{los}} \sim 100$  km s $^{-1}$ , inconsistent with Ret 2 membership. Due to the anomalous nature of this spectrum, we do not include these results in Table 1 or consider them in subsequent analysis.

Finally, the bottom panel of Figure 3 displays a spectrum (Ret2-019) that is difficult to classify given the limited spectral range. Here we consider a few possibilities. First, the absorption features may have molecular origin, specifically from MgH, which has a bandhead at  $\sim 5200$  Å and extends blueward by  $\sim 200$  Å. However, the features in our spectrum are too wide, too deep and too irregularly spaced to match the typical appearance of that band (Weck et al. 2003). In addition, MgH absorption becomes evident typically in stars of spectral types K0 or later, much redder than the de-reddened color ( $g-r \sim 0.17$ ) of this source. Alternatively, this spectrum may reveal the presence of a complex absorption line system along the line of sight to a remote active galactic nucleus. The line widths and extreme depths appear similar to numerous line complexes reported by Srianand et al. (2010) for various absorption lines in Lyman- $\alpha$  systems. Furthermore, the source’s blue color is consistent with that of a background quasi-stellar object (QSO) at moderate redshift. For a number of plausible UV lines that may produce the observed absorption features, the implied redshift of the absorber is  $z \gtrsim 2.5$ , with the background source obviously more distant. The nearest known extragalactic source in the NED catalog is detected in the UV, by GALEX, and is 24 arcsec away from our target. A definitive classification of this source will require spectroscopy with broader wavelength coverage.

## 4. RESULTS

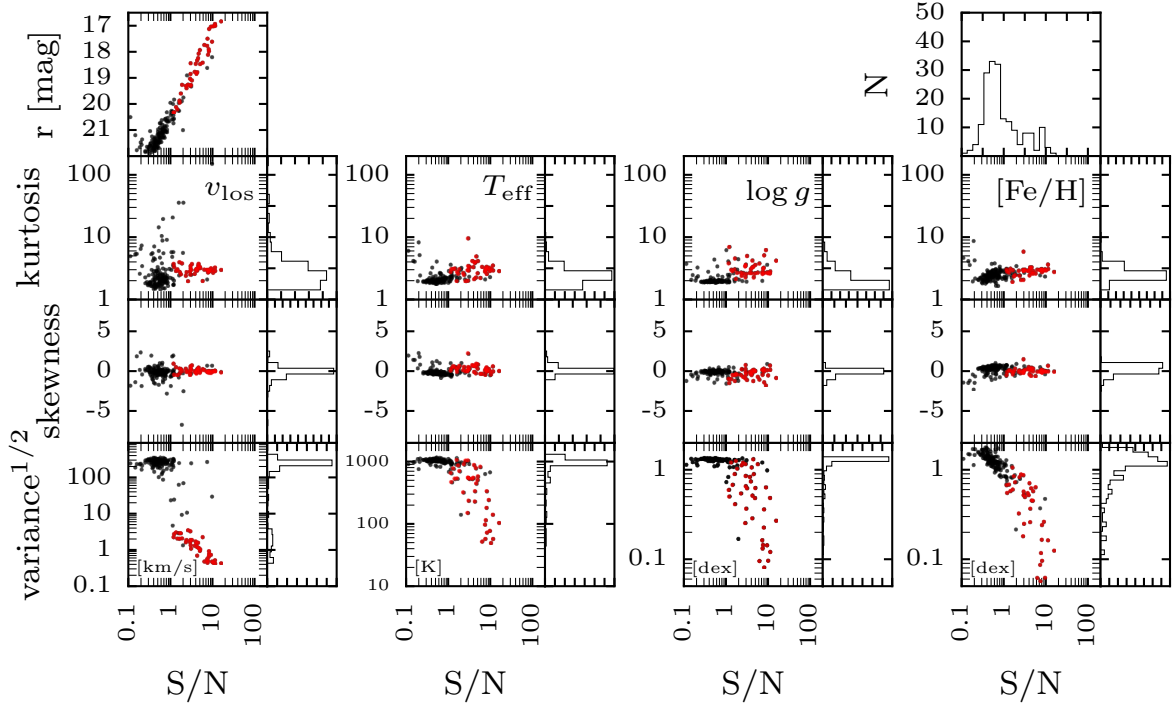


FIG. 4.— Moments of marginalized, 1D posterior PDFs of model parameters vs median S/N per pixel, for all 182 spectra acquired with M2FS. Histograms show marginalized distributions for each moment as well as S/N. Red points indicate observations satisfying quality-control criteria (Section 3.1), where posterior PDFs for velocity estimates are near-Gaussian ( $|\text{skewness}| \leq 1$ ,  $|\text{kurtosis}-3| \leq 1$ ) and the velocity error is  $\sigma \leq 5 \text{ km s}^{-1}$ .

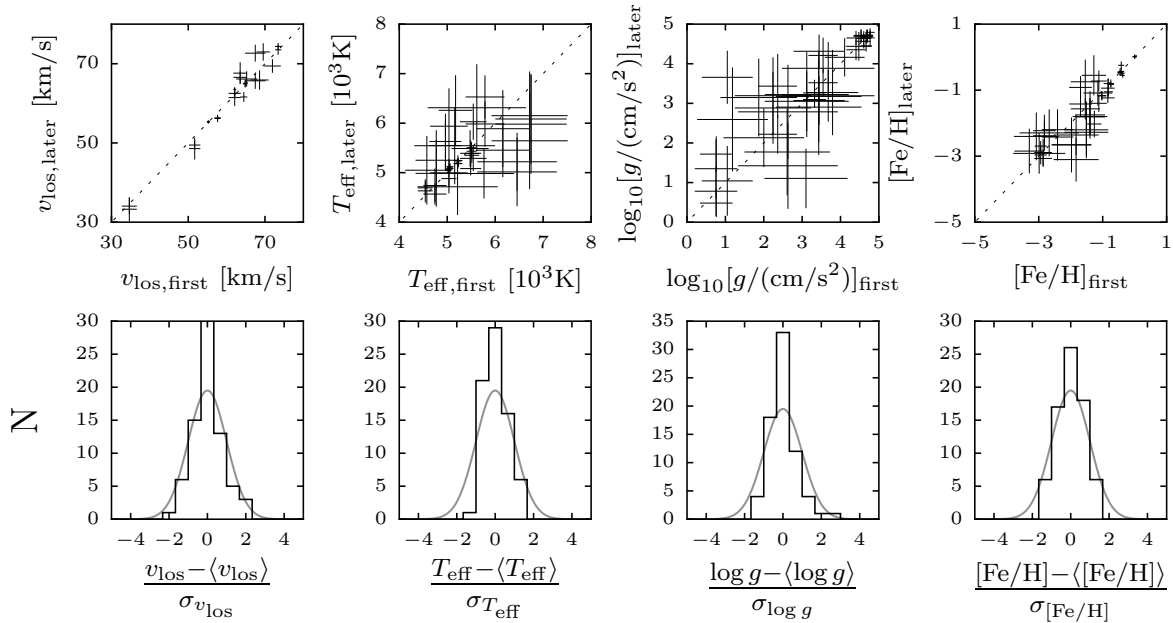


FIG. 5.— Repeatability of parameter estimates, from independent analyses of spectra extracted from three exposures. Scatterplots in the top row display parameter estimates obtained from either of two later exposures against those obtained from the first exposure. Dashed lines are 1:1 relations. Histograms in the bottom row indicate distributions of deviations with respect to the weighted mean for each unique star, normalized by credibility intervals. Overplotted are Gaussian distributions (gray curves) with integrated areas equal to those of the histograms. The empirical distributions are generally narrower than the Gaussian ones, implying that our errors may be slightly overestimated.

TABLE 1  
M2FS STELLAR SPECTROSCOPY OF RETICULUM 2<sup>a</sup>

ID	$\alpha_{2000}$ [hh:mm:ss]	$\delta_{2000}$ [°:':"]	$R$ [arcmin]	$g$ [mag]	$r$ [mag]	S/N <sup>b</sup>	$v_{\text{los}}$ [km s <sup>-1</sup> ] <sup>c</sup>	$T_{\text{eff}}$ [K]	$\log_{10} g$ [dex] <sup>d</sup>	[Fe/H] [dex]	member?
Ret2-004	03:35:56.27	-54:03:16.2	2.1	18.83	18.37	5.7	63.9 ± 2.3 <sup>(0.2,2.0)</sup>	6506 ± 811 <sup>(-0.3,2.3)</sup>	3.12 ± 1.14 <sup>(-0.7,2.7)</sup>	-2.26 ± 0.64 <sup>(-0.6,2.6)</sup>	Y
Ret2-006	03:36:01.76	-54:04:05.4	3.1	19.52	19.05	3.2	62.5 ± 1.9 <sup>(0.7,3.7)</sup>	5313 ± 524 <sup>(0.4,3.3)</sup>	4.02 ± 0.62 <sup>(-1.1,4.9)</sup>	-2.88 ± 0.50 <sup>(0.2,3.0)</sup>	Y
Ret2-016	03:36:35.74	-54:01:56.9	7.9	17.99	17.58	9.5	133.9 ± 0.5 <sup>(0.0,3.0)</sup>	5460 ± 78 <sup>(0.0,2.9)</sup>	4.56 ± 0.14 <sup>(0.1,2.7)</sup>	-0.82 ± 0.06 <sup>(-0.0,3.1)</sup>	N
Ret2-021	03:36:22.83	-53:59:55.5	6.7	18.53	18.07	7.7	128.3 ± 0.5 <sup>(-0.0,2.7)</sup>	5285 ± 115 <sup>(0.3,3.2)</sup>	3.19 ± 0.15 <sup>(0.2,3.1)</sup>	-0.99 ± 0.12 <sup>(0.2,3.1)</sup>	N
Ret2-023	03:36:21.86	-54:00:40.6	6.3	20.23	19.81	1.6	66.7 ± 2.0 <sup>(0.3,4.0)</sup>	5658 ± 1034 <sup>(0.2,2.0)</sup>	3.37 ± 0.89 <sup>(-0.5,3.1)</sup>	-1.96 ± 1.05 <sup>(-0.1,2.0)</sup>	Y
Ret2-032	03:35:40.69	-54:10:05.1	7.1	18.19	17.90	4.9	49.0 ± 1.0 <sup>(-0.2,3.4)</sup>	5786 ± 503 <sup>(0.3,2.7)</sup>	3.81 ± 0.48 <sup>(0.0,2.7)</sup>	-1.75 ± 0.45 <sup>(0.0,2.6)</sup>	N
Ret2-035	03:35:37.06	-54:04:01.2	1.3	18.54	18.03	4.6	62.9 ± 1.2 <sup>(0.4,3.8)</sup>	5034 ± 441 <sup>(0.9,4.9)</sup>	1.50 ± 0.89 <sup>(0.2,2.3)</sup>	-2.90 ± 0.45 <sup>(0.5,3.3)</sup>	Y
Ret2-040	03:35:58.14	-54:02:04.7	2.5	19.25	18.80	4.1	69.1 ± 1.5 <sup>(-0.3,3.1)</sup>	6283 ± 865 <sup>(-0.1,2.2)</sup>	2.97 ± 1.05 <sup>(-0.4,2.7)</sup>	-2.28 ± 0.71 <sup>(-0.3,2.5)</sup>	Y
Ret2-049	03:35:20.36	-54:18:16.9	15.7	20.55	20.26	1.2	308.6 ± 2.3 <sup>(0.9,3.7)</sup>	6009 ± 645 <sup>(0.1,2.8)</sup>	4.06 ± 0.59 <sup>(-0.9,3.8)</sup>	-0.58 ± 0.67 <sup>(-0.4,3.1)</sup>	N
Ret2-062	03:35:09.50	-54:02:29.6	4.8	18.20	17.73	7.3	95.9 ± 0.6 <sup>(-0.1,3.0)</sup>	5426 ± 98 <sup>(0.1,2.8)</sup>	4.49 ± 0.17 <sup>(0.2,2.8)</sup>	-0.87 ± 0.09 <sup>(0.0,3.0)</sup>	N
Ret2-063	03:35:11.65	-54:03:21.8	4.5	19.28	18.80	3.0	70.4 ± 1.9 <sup>(-0.2,3.5)</sup>	4494 ± 548 <sup>(2.1,9.2)</sup>	1.19 ± 0.85 <sup>(0.8,3.5)</sup>	-3.52 ± 0.52 <sup>(1.5,5.7)</sup>	Y
Ret2-069	03:35:02.49	-54:03:54.2	5.9	19.20	18.72	3.7	67.9 ± 1.3 <sup>(-0.0,2.6)</sup>	5048 ± 523 <sup>(0.8,4.1)</sup>	1.90 ± 0.83 <sup>(-0.1,2.7)</sup>	-2.77 ± 0.55 <sup>(0.3,2.9)</sup>	Y
Ret2-072	03:34:37.97	-54:13:59.6	14.5	20.47	20.06	1.4	137.3 ± 3.0 <sup>(0.5,2.9)</sup>	5574 ± 762 <sup>(0.6,3.1)</sup>	3.39 ± 1.20 <sup>(-0.9,3.0)</sup>	-2.13 ± 0.68 <sup>(0.2,2.7)</sup>	N
Ret2-076	03:34:36.69	-54:06:44.9	10.3	18.84	18.30	4.3	71.6 ± 0.7 <sup>(-0.0,2.9)</sup>	4988 ± 229 <sup>(0.7,3.9)</sup>	2.26 ± 0.29 <sup>(0.1,3.3)</sup>	-1.52 ± 0.26 <sup>(0.3,3.3)</sup>	Y
Ret2-077	03:34:53.23	-54:14:03.9	13.2	20.63	20.26	1.2	290.8 ± 2.7 <sup>(-0.4,3.4)</sup>	5372 ± 464 <sup>(0.3,3.1)</sup>	4.28 ± 0.50 <sup>(-1.4,6.8)</sup>	-1.57 ± 0.55 <sup>(-0.0,2.8)</sup>	N
Ret2-079	03:34:54.24	-54:05:58.0	7.6	18.91	18.42	3.9	70.0 ± 1.7 <sup>(-0.2,3.0)</sup>	4915 ± 543 <sup>(0.7,3.3)</sup>	3.44 ± 0.74 <sup>(-0.3,3.1)</sup>	-3.19 ± 0.52 <sup>(0.5,2.9)</sup>	Y
Ret2-080	03:34:47.93	-54:05:25.0	8.3	17.50	16.91	11.2	63.2 ± 0.5 <sup>(0.1,3.3)</sup>	4501 ± 141 <sup>(0.6,3.5)</sup>	0.58 ± 0.48 <sup>(0.7,2.6)</sup>	-3.16 ± 0.16 <sup>(0.7,3.6)</sup>	Y
Ret2-087	03:34:13.05	-53:59:56.1	13.4	19.50	19.00	2.9	200.5 ± 1.7 <sup>(-0.4,2.9)</sup>	5211 ± 150 <sup>(0.0,3.0)</sup>	4.55 ± 0.26 <sup>(-0.8,3.4)</sup>	-1.32 ± 0.18 <sup>(0.1,3.0)</sup>	N
Ret2-089	03:34:05.48	-54:03:49.9	14.2	17.61	17.01	10.1	53.2 ± 0.5 <sup>(-0.0,2.7)</sup>	5107 ± 50 <sup>(-0.4,3.1)</sup>	4.77 ± 0.10 <sup>(-0.5,2.7)</sup>	-0.49 ± 0.04 <sup>(0.0,3.0)</sup>	N
Ret2-090	03:33:41.70	-54:00:07.2	17.9	18.30	18.21	4.5	32.4 ± 1.6 <sup>(0.3,3.5)</sup>	6465 ± 773 <sup>(-0.2,2.3)</sup>	3.43 ± 1.30 <sup>(-1.2,3.4)</sup>	-1.74 ± 0.62 <sup>(-0.5,2.7)</sup>	N
Ret2-092	03:34:49.20	-53:50:19.7	14.8	17.55	16.98	11.3	16.2 ± 0.4 <sup>(0.0,2.9)</sup>	5218 ± 58 <sup>(0.0,2.9)</sup>	4.60 ± 0.12 <sup>(0.2,2.8)</sup>	-0.01 ± 0.05 <sup>(0.0,3.1)</sup>	N
Ret2-100	03:34:39.81	-54:00:58.5	9.4	18.18	17.95	8.1	199.4 ± 0.9 <sup>(-0.2,2.9)</sup>	6368 ± 333 <sup>(0.1,2.6)</sup>	4.05 ± 0.36 <sup>(-0.0,2.7)</sup>	-1.07 ± 0.24 <sup>(-0.1,2.8)</sup>	N
Ret2-102	03:34:18.31	-54:10:06.2	14.2	18.06	17.47	7.5	28.1 ± 0.5 <sup>(-0.1,3.0)</sup>	5118 ± 58 <sup>(-0.4,3.0)</sup>	4.70 ± 0.13 <sup>(-0.4,2.6)</sup>	-0.54 ± 0.06 <sup>(-0.0,3.0)</sup>	N
Ret2-103	03:34:13.93	-54:09:34.4	14.5	18.33	17.71	6.3	20.4 ± 0.5 <sup>(0.1,3.0)</sup>	5069 ± 52 <sup>(-0.3,3.3)</sup>	4.79 ± 0.10 <sup>(-1.0,3.8)</sup>	-0.47 ± 0.06 <sup>(-0.0,3.0)</sup>	N
Ret2-110	03:34:29.93	-54:11:11.7	13.4	19.80	19.35	2.4	290.5 ± 1.9 <sup>(-0.1,2.8)</sup>	5797 ± 888 <sup>(0.5,2.5)</sup>	2.12 ± 1.11 <sup>(0.3,2.4)</sup>	-1.69 ± 0.82 <sup>(0.1,2.4)</sup>	N
Ret2-111	03:34:30.06	-54:09:22.2	12.4	18.20	17.88	8.0	270.1 ± 0.7 <sup>(-0.1,2.9)</sup>	5523 ± 182 <sup>(0.5,3.4)</sup>	3.63 ± 0.23 <sup>(0.3,3.2)</sup>	-1.22 ± 0.17 <sup>(0.2,3.1)</sup>	N
Ret2-115	03:35:31.13	-54:01:48.2	2.0	17.57	17.01	9.0	60.4 ± 0.7 <sup>(-0.2,3.2)</sup>	5164 ± 242 <sup>(0.3,3.3)</sup>	2.44 ± 0.63 <sup>(-1.7,6.0)</sup>	-2.62 ± 0.26 <sup>(0.1,2.9)</sup>	Y
Ret2-126	03:35:20.96	-54:03:48.1	3.2	18.91	18.43	5.0	65.6 ± 1.1 <sup>(-0.0,3.2)</sup>	5185 ± 572 <sup>(0.8,4.1)</sup>	2.91 ± 0.62 <sup>(0.0,3.6)</sup>	-2.67 ± 0.57 <sup>(0.3,2.9)</sup>	Y
Ret2-128	03:35:06.56	-54:06:04.3	6.1	19.74	19.26	2.3	160.5 ± 1.6 <sup>(0.3,2.4)</sup>	5466 ± 319 <sup>(0.2,2.7)</sup>	3.99 ± 0.54 <sup>(-0.5,2.7)</sup>	-1.02 ± 0.33 <sup>(-0.0,2.9)</sup>	N
Ret2-129	03:34:57.57	-54:05:31.4	7.0	18.90	18.40	3.0	61.8 ± 1.4 <sup>(0.3,3.2)</sup>	5296 ± 539 <sup>(0.5,3.5)</sup>	2.99 ± 0.65 <sup>(-0.1,3.6)</sup>	-2.04 ± 0.57 <sup>(-0.1,2.9)</sup>	Y
Ret2-134	03:35:15.17	-54:08:43.0	7.0	19.71	19.26	3.0	70.2 ± 3.3 <sup>(0.6,2.5)</sup>	5875 ± 1019 <sup>(0.1,2.0)</sup>	1.94 ± 1.15 <sup>(0.3,2.3)</sup>	-2.41 ± 0.87 <sup>(-0.1,2.1)</sup>	Y
Ret2-136	03:35:14.01	-54:05:58.1	5.1	19.98	19.55	1.8	66.4 ± 2.9 <sup>(0.3,2.6)</sup>	5582 ± 849 <sup>(0.4,2.6)</sup>	2.78 ± 1.25 <sup>(-0.3,2.2)</sup>	-2.14 ± 0.84 <sup>(0.0,2.4)</sup>	Y
Ret2-138	03:35:13.73	-54:04:56.7	4.6	19.67	19.22	1.9	60.1 ± 2.1 <sup>(0.0,3.8)</sup>	5032 ± 694 <sup>(0.7,3.1)</sup>	2.97 ± 1.02 <sup>(-0.4,2.7)</sup>	-2.62 ± 0.72 <sup>(0.3,2.5)</sup>	Y
Ret2-142	03:35:44.18	-54:01:50.0	1.2	20.34	19.89	1.7	65.0 ± 2.9 <sup>(-0.2,2.8)</sup>	6067 ± 529 <sup>(-0.1,2.8)</sup>	4.39 ± 0.41 <sup>(-1.2,4.5)</sup>	-1.24 ± 0.60 <sup>(-0.3,3.0)</sup>	Y
Ret2-147	03:35:36.93	-53:54:45.1	8.2	20.52	20.15	1.5	302.3 ± 2.1 <sup>(-0.5,3.1)</sup>	5335 ± 524 <sup>(0.5,3.2)</sup>	3.71 ± 0.68 <sup>(-0.4,2.9)</sup>	-1.23 ± 0.58 <sup>(0.1,3.1)</sup>	N
Ret2-153	03:35:46.16	-54:07:33.8	4.6	19.21	18.74	4.3	217.9 ± 1.2 <sup>(0.2,3.1)</sup>	5089 ± 333 <sup>(0.6,4.0)</sup>	2.82 ± 0.51 <sup>(-0.2,4.2)</sup>	-2.16 ± 0.37 <sup>(0.1,3.1)</sup>	N
Ret2-159	03:35:49.95	-54:03:21.4	1.2	19.65	19.22	3.1	293.2 ± 1.1 <sup>(0.2,3.0)</sup>	5329 ± 235 <sup>(0.2,2.9)</sup>	4.26 ± 0.37 <sup>(-0.4,2.6)</sup>	-1.37 ± 0.25 <sup>(0.1,2.9)</sup>	N
Ret2-178	03:36:07.75	-54:02:35.5	3.8	17.39	16.81	15.6	61.8 ± 0.4 <sup>(-0.1,3.0)</sup>	4688 ± 105 <sup>(0.1,2.9)</sup>	1.34 ± 0.27 <sup>(-0.8,4.2)</sup>	-2.74 ± 0.13 <sup>(-0.0,2.8)</sup>	Y

<sup>a</sup>This version lists results only for observations satisfying quality-control criteria (Section 3.1). The online version includes results for all targeted stars.

<sup>b</sup>median signal-to-noise ratio per pixel

<sup>c</sup>line-of-sight velocity in the heliocentric rest frame

<sup>d</sup>units of  $g$  are cm s<sup>-2</sup>



TABLE 2  
COORDINATES AND PHOTOMETRY FOR SPECIAL CASES (SECTION 3.5)

ID	$\alpha_{2000}$ [hh:mm:ss]	$\delta_{2000}$ [°:′:″]	$R$ [arcmin]	$g$ [mag]	$r$ [mag]	notes
Ret2-161	03:36:18.68	-53:57:45.1	7.5	17.99	18.17	BHB candidate
Ret2-034	03:35:39.85	-54:04:58.1	2.0	18.28	18.58	BHB candidate
Ret2-097	03:34:37.33	-53:53:53.9	13.1	17.66	17.10	double star
Ret2-019	03:36:22.99	-53:55:07.9	9.9	21.11	20.94	unknown

Scatterplots in Figure 6 show how the spectroscopically-derived quantities ( $v_{\text{los}}$ ,  $T_{\text{eff}}$ ,  $\log g$ ,  $[\text{Fe}/\text{H}]$ ) relate to each other and to photometrically-derived quantities (position, color, magnitude).

#### 4.1. Membership

Figure 6 also helps to distinguish bona fide members of Ret 2 from contaminants in the Galactic foreground. We expect Ret 2 members to exhibit a relatively narrow range in velocity, lower metallicities, lower surface gravities, and to be clustered nearer Ret 2’s center than are their foreground counterparts. Figure 6 clearly shows a population of stars having these characteristics. Red boxes in Figure 6 enclose measurements that cluster near the centers of Ret 2’s distributions for  $v_{\text{los}}$ ,  $\log g$ ,  $[\text{Fe}/\text{H}]$  and position, as determined by eye. We make no selection based on color, magnitude or temperature because the full ranges for these quantities are consistent with Ret 2 membership. We count 18 stars that are inside *all* red boxes. Table 1 lists a membership status of ‘Y’ for these stars, and ‘N’ for stars that lie outside any of the red boxes.

In previous work we have employed an expectation-maximization (EM) algorithm in order to estimate model parameters in the presence of contamination (Walker et al. 2009b). Under the assumption that velocity dispersion is spatially uniform, our EM algorithm provides estimates of membership probabilities for all individual stars. Reassuringly, when we apply the EM algorithm to the Ret 2 data in Table 1, the sum of membership probabilities is 17.8, with each of the stars inside our red boxes receiving membership probability  $> 0.85$ . Furthermore, the EM algorithm’s estimates of means and dispersions of velocity and metallicity distributions agree well with those obtained in the independent analysis that we present in Section 4.2.

#### 4.2. Chemodynamics of Reticulum 2

In order to characterize the internal chemodynamics of Ret 2, we consider a simple model in which the velocities and metallicities of the 18 member stars are normally distributed about means that can vary systematically with position in order to account for velocity and metallicity gradients. Specifically we assume that the mean line-of-sight velocity,  $\mu_{v_{\text{los}}}$ , is a function of projected position ( $R, \theta$ ),

$$\mu_{v_{\text{los}}}(R, \theta) \equiv \langle v_{\text{los}} \rangle + k_{v_{\text{los}}} R \cos(\theta - \theta_{v_{\text{los}}}), \quad (2)$$

where  $\langle v_{\text{los}} \rangle$  is the mean velocity at the center,  $k_{v_{\text{los}}} \equiv |d\mu_{v_{\text{los}}}/dR|$  is the magnitude of maximum velocity gradient and  $\theta_{v_{\text{los}}}$  (measured from north of Ret 2’s center and opening to the east, in equatorial coordinates) specifies its direction. We assume that mean metallicity,  $\mu_{[\text{Fe}/\text{H}]}$ , is a function only of separation from the center,

$$\mu_{[\text{Fe}/\text{H}]}(R) \equiv \langle [\text{Fe}/\text{H}] \rangle + k_{[\text{Fe}/\text{H}]} R, \quad (3)$$

where  $\langle [\text{Fe}/\text{H}] \rangle$  is the mean metallicity at the center and  $k_{[\text{Fe}/\text{H}]} \equiv d\mu_{[\text{Fe}/\text{H}]} / dR$  specifies the magnitude of any (isotropic) metallicity gradient.

Under these assumptions, the joint probability density for observables  $v_{\text{los}}$  and  $[\text{Fe}/\text{H}]$ , for a star at position ( $R, \theta$ ), is

$$p(v_{\text{los}}, [\text{Fe}/\text{H}] | R, \theta) = \frac{(2\pi)^{-1}}{\sqrt{(\sigma_{v_{\text{los}}}^2 + \delta_{v_{\text{los}}}^2)(\sigma_{[\text{Fe}/\text{H}]}^2 + \delta_{[\text{Fe}/\text{H}]}^2)}} \times \exp \left[ -\frac{1}{2} \frac{(v_{\text{los}} - \mu_{v_{\text{los}}})^2}{\sigma_{v_{\text{los}}}^2 + \delta_{v_{\text{los}}}^2} - \frac{1}{2} \frac{([\text{Fe}/\text{H}] - \mu_{[\text{Fe}/\text{H}]})^2}{\sigma_{[\text{Fe}/\text{H}]}^2 + \delta_{[\text{Fe}/\text{H}]}^2} \right] \quad (4)$$

where  $\sigma_{v_{\text{los}}}$  and  $\sigma_{[\text{Fe}/\text{H}]}$  are velocity and metallicity dispersions, respectively, and  $\delta_{v_{\text{los}}}$  and  $\delta_{[\text{Fe}/\text{H}]}$  are measurement errors. Given a vector of free parameters  $\vec{\theta} \equiv (\langle v_{\text{los}} \rangle, \sigma_{v_{\text{los}}}, k_{v_{\text{los}}}, \theta_{v_{\text{los}}}, \langle [\text{Fe}/\text{H}] \rangle, \sigma_{[\text{Fe}/\text{H}]}, k_{[\text{Fe}/\text{H}]})$ , a data set consisting of  $N$  observations of Ret 2 members,  $D \equiv \{(R_i, \theta_i, v_{\text{los},i}, [\text{Fe}/\text{H}]_i)\}_{i=1}^N$ , has likelihood

$$\mathcal{L}(D | \vec{\theta}) = \prod_{i=1}^N p(v_{\text{los},i}, [\text{Fe}/\text{H}]_i | R_i, \theta_i). \quad (5)$$

From Bayes’ theorem, given the data, the model has posterior probability distribution function (PDF)

$$p(\vec{\theta} | D) = \frac{\mathcal{L}(D | \vec{\theta}) p(\vec{\theta})}{p(D)}, \quad (6)$$

where  $p(\vec{\theta})$  is the prior PDF and  $p(D) \equiv \int \mathcal{L}(D | \vec{\theta}) p(\vec{\theta}) d\vec{\theta}$  is the ‘evidence’.

Again we use MultiNest to scan the 7-dimensional parameter space and to draw random samples from the posterior PDF. Figure 7 displays random samplings from PDFs for the means and dispersions. Table 3 summarizes the results. For each of the seven free parameters, the second column lists boundaries inside which the prior is uniform and nonzero (outside these boundaries, the prior probability is zero). The third column lists median-likelihood values and credibility intervals that enclose the central 68% (95%) of area under the posterior PDF.

We find that Ret 2 has mean line-of-sight velocity  $\langle v_{\text{los}} \rangle = 64.8^{+1.1}_{-1.0} \text{ km s}^{-1}$ , with a resolved velocity dispersion of  $\sigma_{v_{\text{los}}} = 3.6^{+0.9}_{-0.6} \text{ km s}^{-1}$ . We obtain a mean metallicity of  $\langle [\text{Fe}/\text{H}] \rangle = -2.67^{+0.34}_{-0.34}$ , with a dispersion of  $\sigma_{[\text{Fe}/\text{H}]} = 0.50^{+0.17}_{-0.13}$ . Taken at face value, the estimated metallicity dispersion is five times larger than is exhibited by any known globular cluster (Willman & Strader 2012). However, we note that the estimated metallicity dispersion is also similar to the median credibility interval for our metallicity estimates. Koposov et al. (2011) demonstrate that, in such cases, estimates of intrinsic dispersion can be extremely sensitive to systematic over- or underestimation of measurement errors. Given the behavior in the bottom-right panel of Figure 5 (see Section 3.3 for discussion), we suspect that our metallicity errors are systematically overestimated, which would imply that our estimate of Ret 2’s intrinsic metallicity dispersion is underestimated. In any case, this estimate should be interpreted with caution and will need to be confirmed using spectra spanning a larger range of wavelengths.

Finally, we do not detect significant gradients in either velocity or metallicity. Our estimate of  $k_{v_{\text{los}}} = 0.4^{+0.4}_{-0.2} \text{ km s}^{-1}$

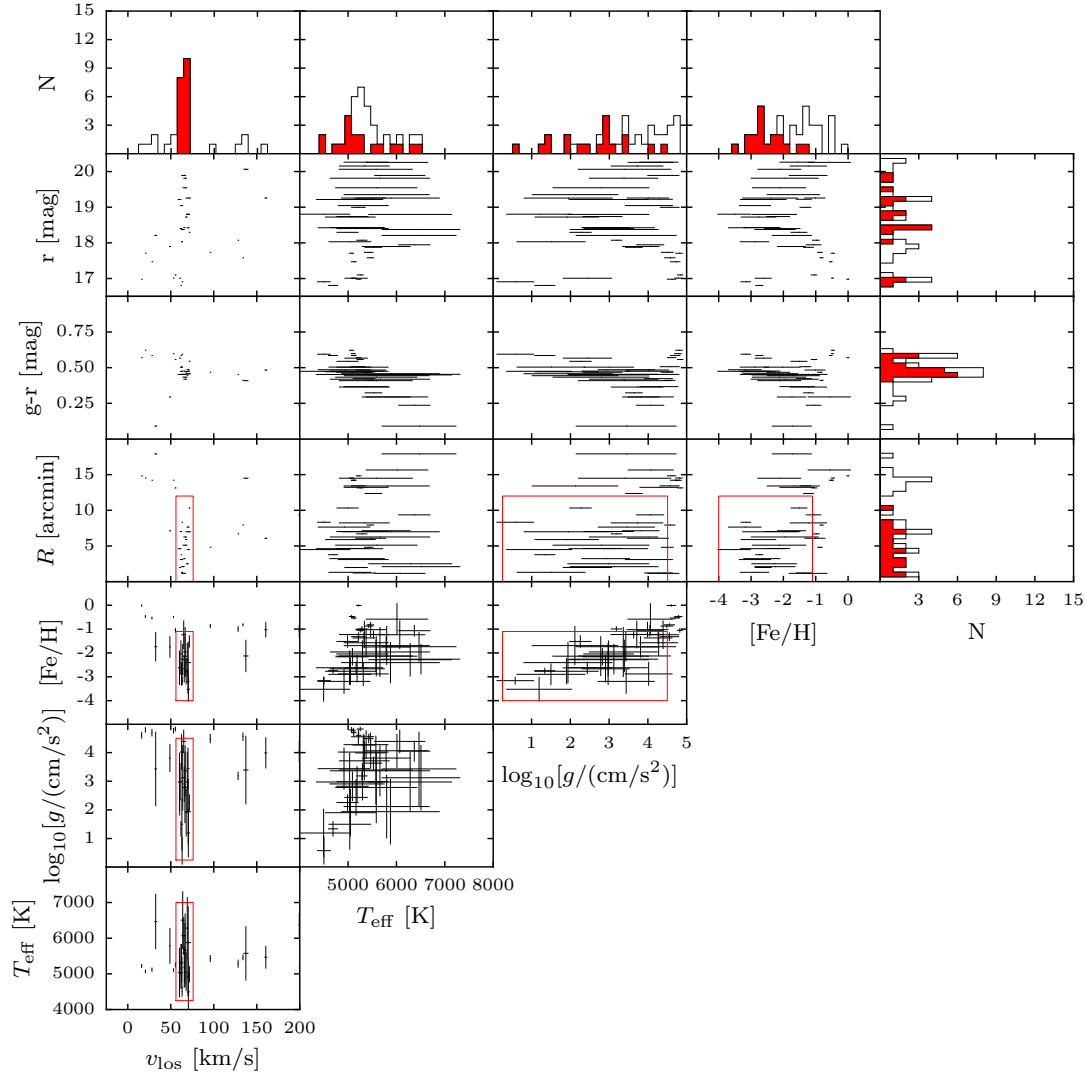


FIG. 6.— Scatterplots showing relations among photometrically- (magnitude,  $r$ , color,  $g-r$ , and separation,  $R$ , from Ret 2's center) and spectroscopically-derived ( $v_{\text{los}}$ ,  $T_{\text{eff}}$ ,  $\log g$  and  $[\text{Fe}/\text{H}]$ ) quantities for individual Ret 2 stars. Red boxes are drawn, by eye, to enclose probable members. For each observable, histograms show 1D distributions for the full sample (open black) and likely members (solid red).

TABLE 3  
SUMMARY OF PROBABILITY DISTRIBUTION FUNCTIONS FOR CHEMODYNAMICAL PARAMETERS

parameter	prior	posterior	description
$\langle v_{\text{los}} \rangle$ [ $\text{km s}^{-1}$ ]	uniform between -500 and +500	$64.8^{+1.1(+2.2)}_{-1.0(-2.0)}$	mean velocity at center
$\sigma_{v_{\text{los}}}$ [ $\text{km s}^{-1}$ ]	uniform between 0 and +500	$3.6^{+0.9(+2.1)}_{-0.6(-1.1)}$	velocity dispersion
$k_{v_{\text{los}}}$ [ $\text{km s}^{-1} \text{ arcmin}^{-1}$ ]	uniform between 0 and +10	$0.4^{+0.4(+1.0)}_{-0.2(-0.3)}$	magnitude of maximum velocity gradient
$\theta_{v_{\text{los}}}$ [ $^{\circ}$ ]	uniform between -180 and +180	$-79^{+220(+251)}_{-55(-93)}$	direction of maximum velocity gradient
$\langle [\text{Fe}/\text{H}] \rangle$	uniform between -5 and +1	$-2.67^{+0.34(+0.72)}_{-0.34(-0.67)}$	mean metallicity at center
$\sigma_{[\text{Fe}/\text{H}]}$	uniform between 0 and +2	$0.50^{+0.17(+0.40)}_{-0.13(-0.23)}$	metallicity dispersion
$k_{[\text{Fe}/\text{H}]}$ [ $\text{dex arcmin}^{-1}$ ]	uniform between -1 and +1	$0.02^{+0.06(+0.12)}_{-0.06(-0.13)}$	magnitude of metallicity gradient

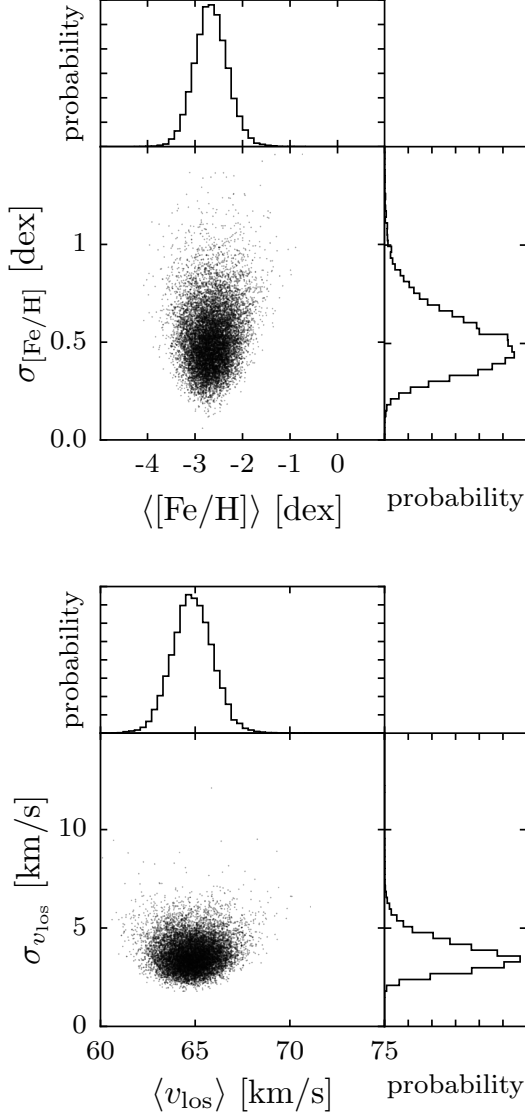


FIG. 7.— Samples drawn randomly from posterior PDFs for means and dispersions of Ret 2’s metallicity (top) and velocity (bottom) distributions. Histograms display marginalized, 1D PDFs for each parameter.

$\text{arcmin}^{-1}$  is consistent with zero at the  $\lesssim 2\sigma$  level. On the other hand, values as large as  $\sim 1 \text{ km s}^{-1} \text{ arcmin}^{-1}$  are similarly allowed. The steepest allowed gradients are directed along position angle  $\theta_{v_{\text{los}}} \sim -80$  degrees, roughly  $\sim 30$  degrees from Ret 2’s morphological major axis (K15, DES15). Our estimate of  $k_{[\text{Fe}/\text{H}]} = 0.02^{+0.06}_{-0.06} \text{ dex arcmin}^{-1}$  is consistent with zero at the  $\sim 1\sigma$  level, but also allows gradients as steep as  $\pm 0.15 \text{ dex arcmin}^{-1}$  within  $\sim 2\sigma$ . We notice degeneracies between our estimates of dispersions and gradients for both velocity and metallicity distributions, such that larger gradients correspond to smaller dispersions. However, our estimates for individual parameters (Table 3) naturally include such effects, as the 1D PDFs for individual parameters are obtained by marginalizing over all other dimensions of the parameter space.

#### 4.3. Scaling Relations

K15 use photometric data to show that Ret 2 occupies a region of structural parameter space that is populated by objects of ambiguous classification, somewhat intermediate between well-established globular clusters and dwarf galaxies (Gilmore et al. 2007, cf. K15’s Figure 17). Specifically, with projected half-light radius  $R_h = 32 \pm 1 \text{ pc}$ , Ret 2 is larger than nearly all globular clusters and smaller than nearly all known galaxies. Moreover, its absolute magnitude,  $M_V = -2.7 \pm 0.1$ , would place Ret 2 among the least luminous members of either population. In these regards, Ret 2 and many of its newly-discovered siblings are similar to ‘ultrafaint’ satellites Segue 1, Segue 2 and Willman 1, as well as to the extended globular clusters Pal 14 and Crater (Belokurov et al. 2014; Laevens et al. 2014, Mateo et al. 2015).

Our spectroscopic results provide new leverage that can settle the question of Ret 2’s nature. For Galactic globular clusters as well as the dwarf spheroidal satellites of the Milky Way and M31, the top panel of Figure 8 plots mean metallicity against luminosity. While globular clusters show no obvious trend, it is well-known that dwarf galaxies follow a luminosity/metallicity relation (Mateo 1998; Tolstoy et al. 2009; Kirby et al. 2013). Given the low mean metallicity that we estimate from the M2FS spectra,  $\langle [\text{Fe}/\text{H}] \rangle = -2.67^{+0.34}_{-0.34}$ , we place Ret 2 squarely onto the galactic relation (large black square in Figure 8).

The lower panel of Figure 8 shows another well-known scaling relation that distinguishes dwarf galaxies from globular clusters. Specifically, the mass-to-light ratios of dwarf galaxies are anti-correlated with luminosity, such that gravitational potentials in the least luminous galaxies all seem to be dominated by dark matter (Mateo et al. 1993; Walker et al. 2007; Martin et al. 2007; Simon & Geha 2007). In contrast, the stellar kinematics of globular clusters generally do not require an internal dark matter component. In order to see these trends, we plot  $R_h \sigma_{v_{\text{los}}}^2 / (GL_V)$  against luminosity, where  $R_h$  is half-light radius,  $\sigma_{v_{\text{los}}}$  is line-of-sight velocity dispersion,  $L_V$  is total V-band luminosity and  $G$  is Newton’s constant. The combination of macroscopic observables on the vertical axis has dimensions of mass-to-light ratio<sup>5</sup> and is therefore sufficient to highlight the different behavior of dwarf galaxies and globular clusters. Again we find that Ret 2 follows the galactic relation. Moreover, assuming dynamic equilibrium and negligible contamination from binary stars, Ret 2 has among the highest dynamical mass-to-light ratios of any known object. The crude mass estimator of Walker et al. (2009a) implies that the dynamical mass enclosed within Ret 2’s projected half-light radius is  $M(R_h) \approx 5R_h \sigma_{v_{\text{los}}}^2 / (2G) = 2.4^{+1.3}_{-0.8} \times 10^5 M_\odot$ , and the associated mass-to-light ratio is  $\approx 2M(R_h)/L_V = 462^{+264}_{-157}$  in solar units.

Table 4 summarizes the observed properties of Ret 2, based on the photometric results of K15 and the spectroscopic results presented in this work. Where photometric results from DES15 differ from those of K15 (e.g., for absolute magnitude), we list the DES15 results as well.

#### 5. SUMMARY AND DISCUSSION

We have presented results from our initial stellar-spectroscopic observations of Ret 2. Integrating for two hours in below-average observing conditions, Magellan/M2FS has

<sup>5</sup> Many popular dynamical mass estimators have  $M_{\text{dyn}} \propto R_h \sigma_{v_{\text{los}}}^2 / G$ , where the constant of proportionality is typically between  $\sim 2$ – $6$  (e.g., Richstone & Tremaine 1986; Wolf et al. 2010; Amorisco & Evans 2011).

TABLE 4  
SUMMARY OF RETICULUM 2’S OBSERVED PHOTOMETRIC AND SPECTROSCOPIC PROPERTIES

Quantity	value	reference	notes
R.A. at center	$\alpha_{J2000} = 03:35:42$	K15 <sup>†</sup>	
Dec. at center	$\delta_{J2000} = -54:02:57$	K15	
Galactic longitude	$l = 266.2958$ deg	K15	
Galactic latitude	$b = -49.7357$ deg	K15	
distance modulus	$m - M = 17.4 \pm 0.2$	K15	
distance from Sun	$D \sim 30$ kpc	K15	
absolute magnitude	$M_V = -2.7 \pm 0.1$ ( $-3.6 \pm 0.1$ )	K15 (DES15)	
exponential scale length	$R_e = 3.37^{+0.23}_{-0.13}$ arcmin	K15	semi-major axis
ellipticity	$e = 1 - (b/a) = 0.59^{+0.02}_{-0.03}$	K15	
position angle	$PA = 71 \pm 1$ deg	K15	
projected halfight radius	$R_h = 3.64^{+0.21}_{-0.12}$ arcmin	K15	$R_h \approx 1.68 R_e \sqrt{1-e}$
projected halfight radius	$R_h = 32^{+1.9}_{-1.1}$ pc	K15	
systemic line-of-sight velocity	$v_{\text{los}} = 64.8^{+1.1}_{-1.0}$ km s <sup>-1</sup>	this work	solar rest frame
systemic line-of-sight velocity	$v_{\text{los}} = -90.9$ km s <sup>-1</sup>	this work	Galactic rest frame, given solar motion measured by Schönrich et al. (2010)
internal velocity dispersion	$\sigma_{v_{\text{los}}} = 3.6^{+0.9}_{-0.6}$ km s <sup>-1</sup>	this work	
velocity gradient	$k_{v_{\text{los}}} = 0.4^{+0.4}_{-0.2}$ km s <sup>-1</sup> arcmin <sup>-1</sup>	this work	
PA of velocity gradient	$\theta_{v_{\text{los}}} = -79^{+220}_{-55}$ deg	this work	
mean metallicity	$\langle [\text{Fe}/\text{H}] \rangle = -2.67^{+0.34}_{-0.34}$ dex	this work	
metallicity dispersion	$\sigma_{[\text{Fe}/\text{H}]} = 0.50^{+0.17}_{-0.13}$ dex	this work	similar to median metallicity error
metallicity gradient	$k_{[\text{Fe}/\text{H}]} = 0.02^{+0.06}_{-0.06}$ dex arcmin <sup>-1</sup>	this work	
mass enclosed within $R_h$	$M(R_h) = 2.4^{+1.3}_{-0.8} \times 10^5 M_\odot$	this work	$M(R_h) \approx 5R_h \sigma_{v_{\text{los}}}^2 / (2G)$ ; assumes equilibrium, negligible binary stars
mass-to-light raio	$\Upsilon = 462^{+264}_{-157} M_\odot / L_\odot$	this work	$\Upsilon \approx 2M(R_h) / L_V$ ; assumes equilibrium, negligible binary stars

<sup>†</sup> Unless otherwise noted, K15 and DES15 report similar values.

delivered spectra sufficient for estimating  $v_{\text{los}}$ ,  $T_{\text{eff}}$ ,  $\log g$  and  $[\text{Fe}/\text{H}]$  for 38 stars, reaching a limiting magnitude of  $r \lesssim 20.3$ . Based on a combination of photometrically- and spectroscopically-derived quantities, we have confirmed 18 likely members of Ret 2. From the member sample we estimate a velocity dispersion of  $3.6^{+0.9}_{-0.6}$  km s<sup>-1</sup> about a mean of  $64.8^{+1.1}_{-1.0}$  km s<sup>-1</sup> in the solar rest frame, and a metallicity dispersion of  $0.50^{+0.17}_{-0.13}$  dex about a mean of  $-2.67^{+0.34}_{-0.34}$  dex. While Ret 2’s size and luminosity are similar to those of globular clusters as well as ultrafaint dwarf galaxies, our estimates for mean metallicity and velocity dispersion place Ret 2 on scaling relations that are followed by dwarf galaxies and not by globular clusters. On this basis, we conclude that Ret 2 is a bona fide galaxy.

Furthermore, under assumptions of dynamic equilibrium and negligible contamination from binary stars, Ret 2’s combination of size, velocity dispersion and luminosity implies a dynamical mass-to-light ratio of  $\sim 462^{+264}_{-157} M_\odot / L_\odot$ , among the largest inferred for any known object. However, the validity of these assumptions remains unclear. Perhaps undermining the assumption of dynamic equilibrium is Ret 2’s significantly flattened morphology (K15, DES15), which—given Ret 2’s Galactocentric distance of  $\sim 30$  kpc—may signal ongoing tidal disruption. On the other hand, our velocity sample does not show a significant velocity gradient that might reflect the ordered motions associated with tidal streaming (Piatek & Pryor 1995; Read et al. 2006). The mean velocity that we estimate for Ret 2 becomes  $v_{\text{los}} \sim -90.9$  km s<sup>-1</sup> in the Galactic rest frame, implying that Ret 2 is currently approaching pericenter. We expect that deeper photometric and deeper/wider spectroscopic followup will be required in order to determine the extent to which Ret 2 is undergoing a disruptive tidal encounter with the Milky Way.

Finally, even if Ret 2 is resilient to tides, its velocity dispersion reflects some combination of its internal gravitational potential and the intrinsic velocity variability of binary stars. McConnachie & Côté (2010) have demonstrated that binary

motions alone can contribute velocity dispersions of up to a few km s<sup>-1</sup>. Therefore we cannot exclude the possibility that Ret 2’s velocity dispersion receives a significant contribution from binary stars, in which case the dynamical mass and mass-to-light ratios that we have reported would be systematically overestimated. More epochs of spectroscopic observations of Ret 2—and other ultrafaint dwarfs with small velocity dispersions—are required in order to gauge the magnitude of this effect.

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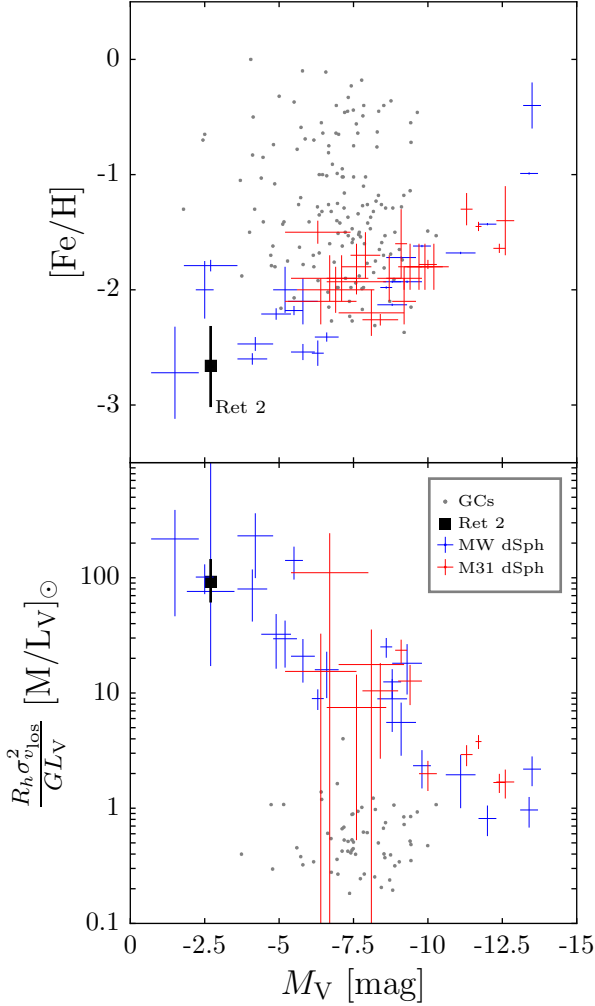


FIG. 8.— Mean metallicity and dynamical mass-to-light ratio vs absolute magnitude, for Galactic globular clusters (black points) as well as dwarf spheroidal satellites of the Milky Way (blue points with errorbars) and M31 (red points with errorbars). Quantities plotted for Ret 2 are adopted from K15 and this work. Data for globular clusters and dSphs are adopted, respectively, from the catalog of Harris (1996, 2010 edition) and the review of McConnachie (2012).

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